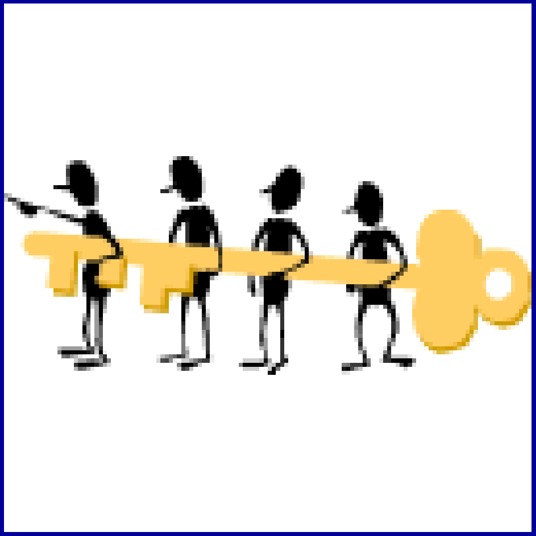
TOPIC # 6
The RADIATION LAWS

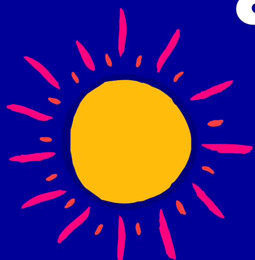
More KEYS
to unlocking the topics of:
The GREENHOUSE EFFECT,
GLOBAL WARMING &
OZONE DEPLETION!

Topic #6
pp 33-38



OBJECTIVES FOR TODAY'S CLASS:

To understand the essentials about the key differences between



Solar radiation

&

Terrestrial radiation

based on the principles of the **"Radiation Laws."**

Another "cartoon" view of Solar vs Terrestrial radiation:

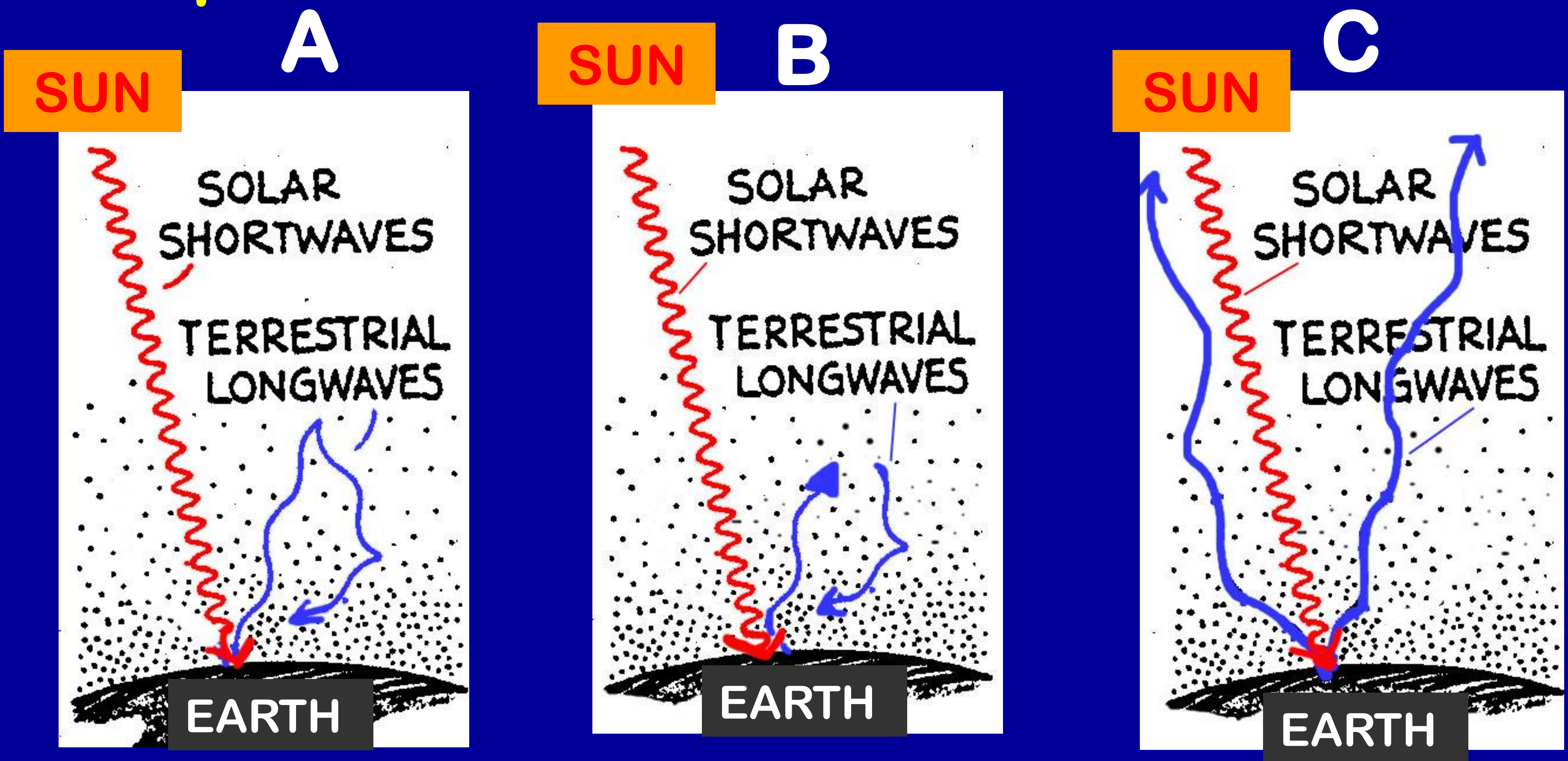
Both Sun & Earth are radiating energy . . . .

NOT TO SCALE!!!

Fire up your clickers . . . . Channel 41

☺

Q1- Which one is the most accurate depiction of the Greenhouse Effect??

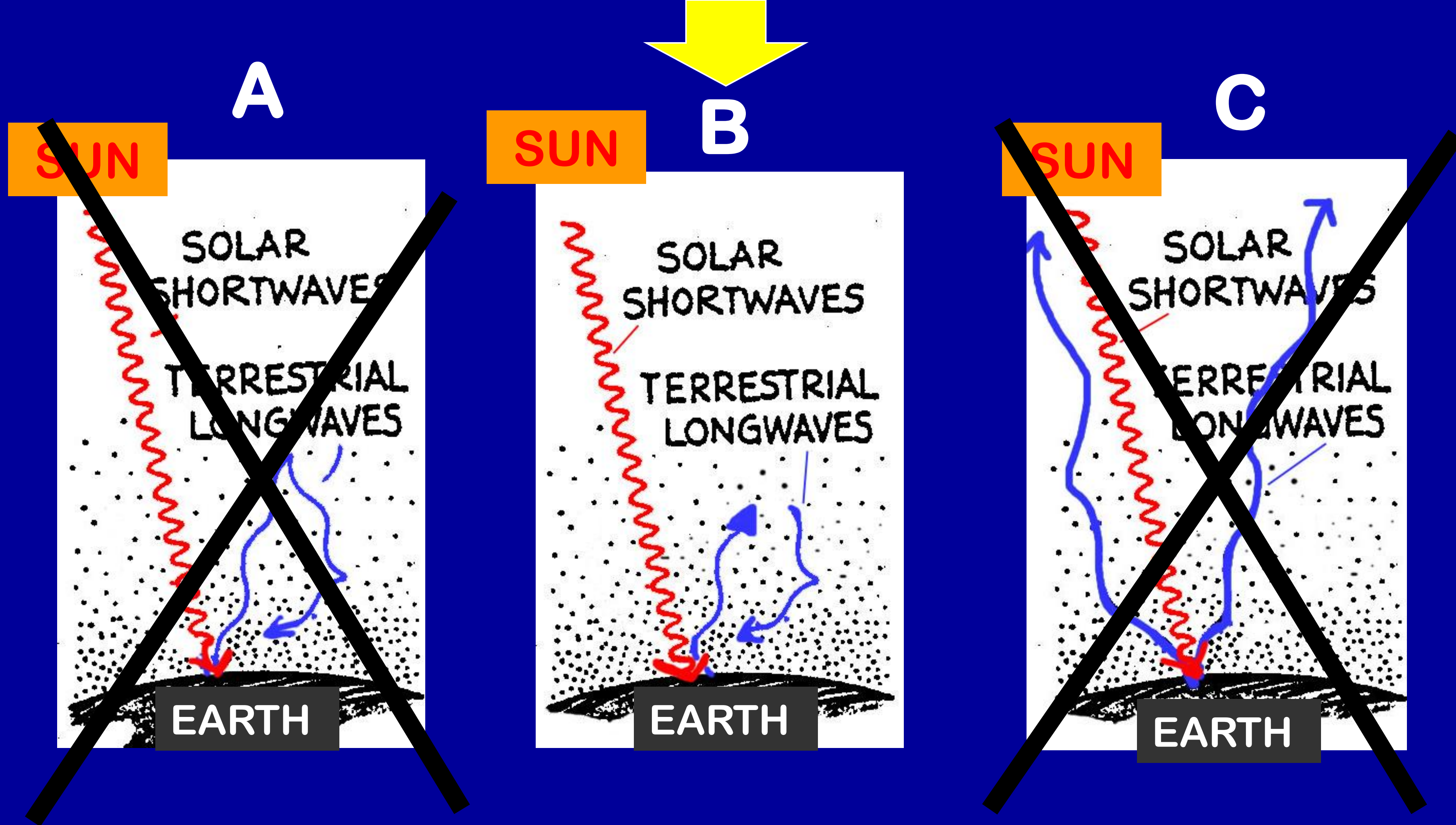


In these & upcoming figures, for convenience:

= **solar (shortwave) radiation** (High Energy)

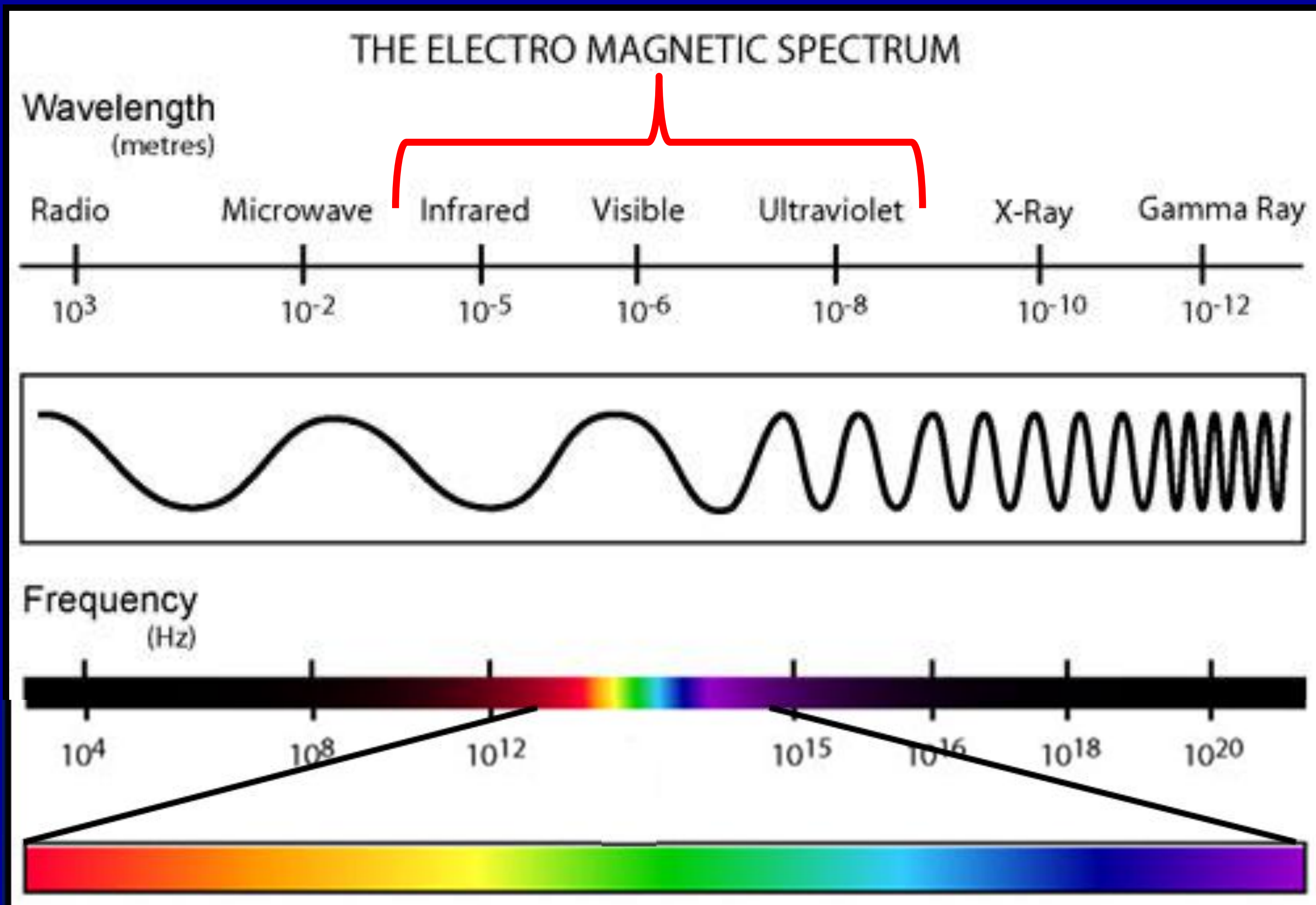
= **terrestrial (longwave) radiation** (Lower Energy)

= gases in the atmosphere



IMPORTANT: Actually, NONE of these is exactly correct, and we will learn why soon. . . .

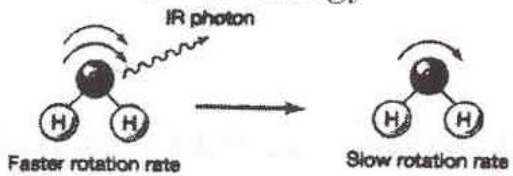
but B is preferred for now.


THE ELECTRO MAGNETIC SPECTRUM
Wavelength
(metres)
Radio
Microwave
Infrared
Visible
Ultraviolet
X-Ray
Gamma Ray
10^{3}
10^{-2}
10^{-5}
10^{-6}
10^{-8}
10^{-10}
10^{-12}
Frequency
(Hz)
10^{4}
10^{8}
10^{12}
10^{15}
10^{16}
10^{18}
10^{20}

R O Y G B (I) V

KEY BANDS IN THE SPECTRUM FOR GLOBAL CHANGE:

UV, Visible, IR, NIR,

Type of Electromagnetic Radiation	Range of Wavelengths (in units indicated)	Typical Source
Gamma rays	10^{-16} to 10^{-11} in meters (m) using scientific notation	high-energy processes within nucleus caused by the strong force
Ultraviolet radiation	.0001 to 0.4 in micrometers (µm)	electrons moving (quantum leaps) within individual atoms
Visible light	0.4 to 0.7 in micrometers (µm)	
Infrared radiation	0.7 to ~30 (up to 1000) in micrometers (µm)	chaotic thermal kinetic motion of molecules due to their thermal energy
Near Infrared radiation	0.7 - 1.0 in micrometers (µm)	
Far Infrared	1.0 - ~30 (up to 1000) in micrometers (µm)	
Microwaves	10^{-4} to 10^{-2} in meters (m) using scientific notation	electronically produced by microwave oven
AM Radio waves	10 to 10^{2} in meters (m) using scientific notation	electronically produced -- waves vibrate in human-made electrical circuits

Solar SW

Terrestrial LW (IR)

GREENHOUSE EFFECT

Review p 32

$E = \sigma T^4$

$(1/d^2)$

"The equations we seek are the poetry of nature . . . . Why is nature that way?

Why is it possible for these powerful manifestations of forces to be trapped in a very simple, beautiful formula?"

"This has been a question which many people have discussed, but there is no answer."

~ Chen Ning Yang (b. 1922) US physicist

$E = hc/\lambda$

$\lambda_m = a/T$

Presenting . . . .

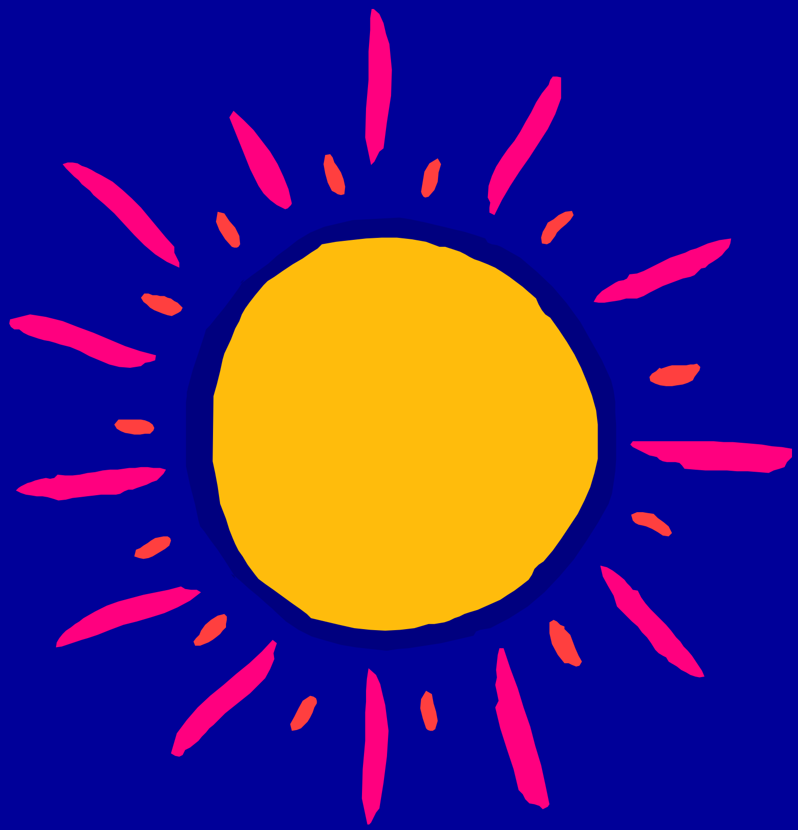
THE RADIATION LAWS !!!

Keys to Understanding the Greenhouse Effect

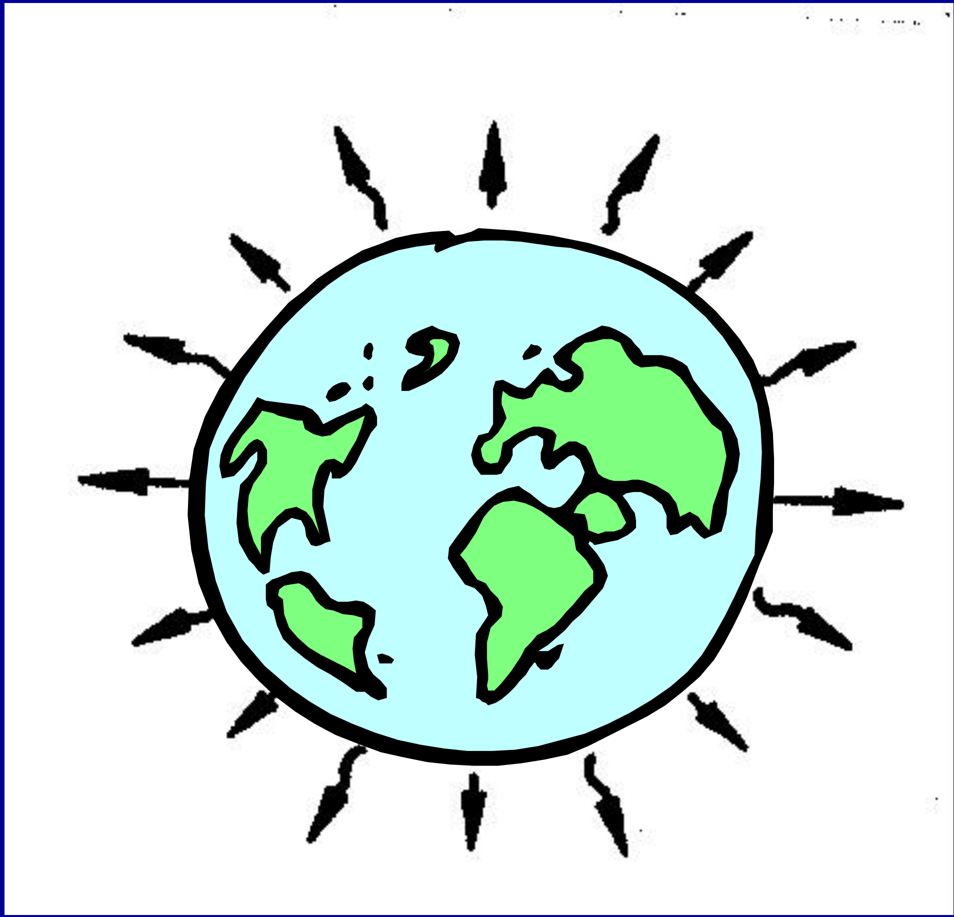
TYING IT ALL TOGETHER: THE RADIATION "LAWS"

The Sun's energy (solar) is emitted in the form of electromagnetic radiation.

mostly SW (but also some LW)



The Earth's energy (terrestrial) is also emitted in the form of electromagnetic wavelengths.



mostly LW

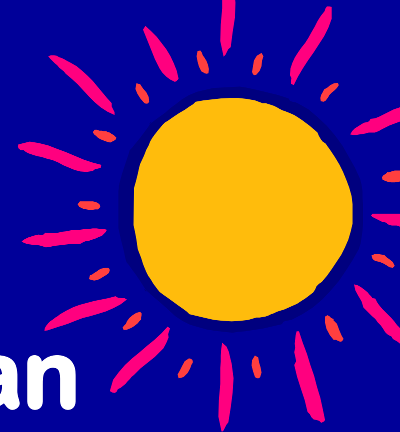
LAW #1

Emission of radiation

All substances emit radiation as long as their temperature is above absolute zero

(-273.15°C or 0 Kelvin).

LAW #2 BLACKBODY & PLANCK FUNCTION CONCEPT



The Sun is very similar to an "ideal emitter" (or "Black body")

(NOTE: the Earth is NOT as good a "black body" as the Sun)

Black body (def): a body that **emits** equally well at **all wavelengths**

(i.e. radiates with 100% efficiency)

It also **absorbs** equally well at all wavelengths and is a **"perfect absorber"**

(hence described as "black")

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Blackbodies ("ideal emitters") exhibit a *defined relationship* between:

the <u>intensity</u> of radiation <u>energy</u> (E)
(i.e. amount of radiation flux) they give off
&
the <u>wavelength</u> of that radiation.

This relationship is called the <u>Planck function</u> :

E = h * speed of light / wavelength

Where (h) is Planck's constant.

or

$$E = h\,c\,/\,\lambda$$

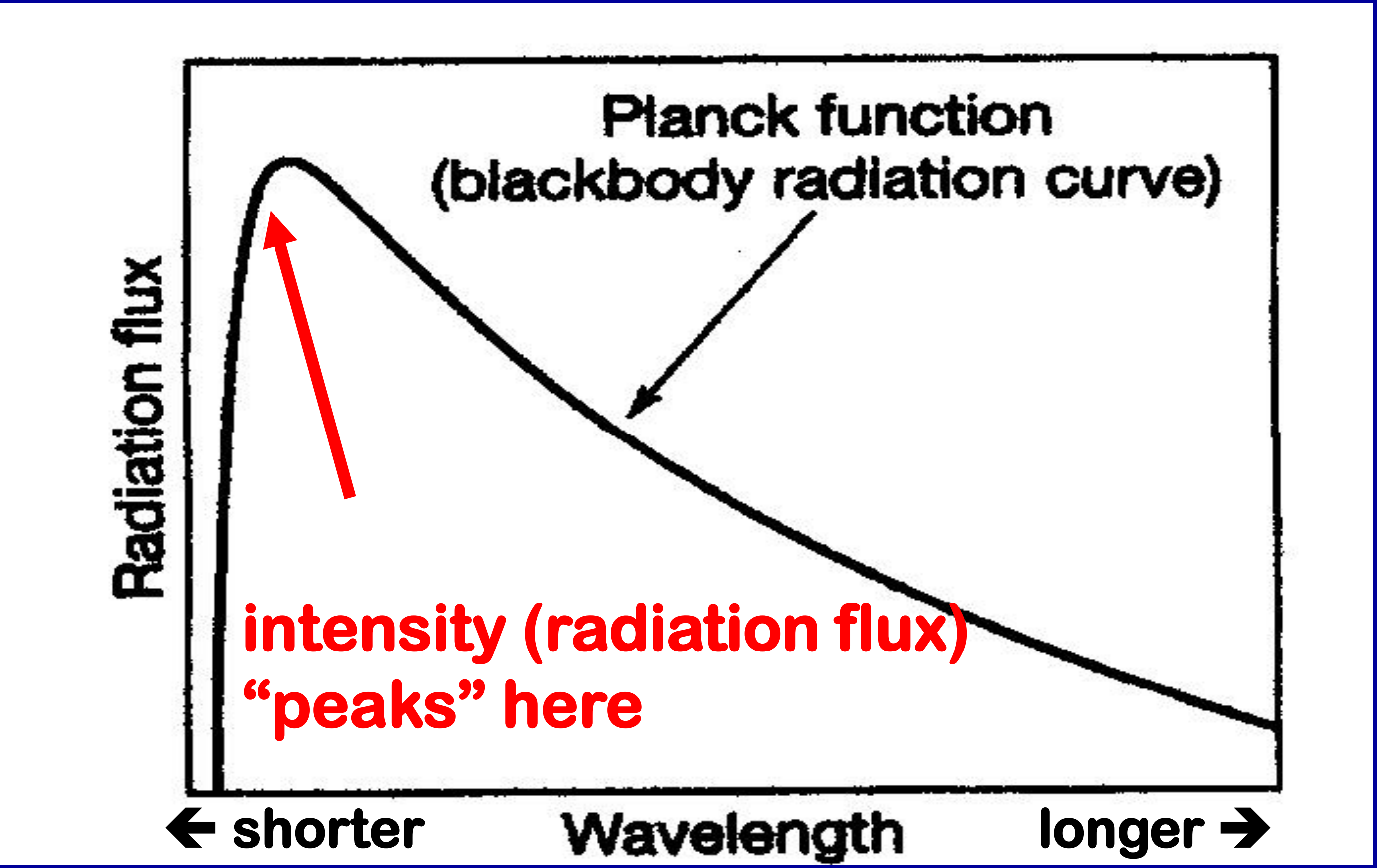
Planck Function:

The Sun emits energy at ALL wavelengths . . .

but the <u>amount</u> of <u>Energy</u> emitted is inversely related to the <u>wavelength of emission</u>

“I radiate at the speed of light like a blackbody; but my energy flux is GREATEST at SHORTER wavelengths”

This can be depicted in a graph:



An emitting blackbody's SHORTER wavelengths have HIGHER intensity radiation (and greater energy flux) than the LONGER wavelengths

Easy way to remember the PLANCK FUNCTION / BLACKBODY concept:

"The shorter the wavelength, the GREATER the intensity of the energy flux"

☺

Q2 - Gamma radiation involves a <u>greater</u> energy flux than microwave radiation.

1. True
2. False
3. Both wavelength bands have the same energy flux
4. We haven’t learned enough yet to answer this!

Q2 - Gamma radiation involves a greater energy flux than microwave radiation.

1. True
2. False
3. Both wavelength bands have the same energy flux
4. We haven’t learned enough yet to answer this!

LAW #3: THE STEFAN-BOLTZMANN LAW:

If the substance is an ideal emitter (black body),

The total **AMOUNT** of radiation given off is <u>proportional</u> to the fourth power of its absolute **TEMPERATURE.**

$E = \sigma T^4$

where σ is a constant

(the Stefan-Boltzmann constant) which has a value of

5.67×10^{-8} W/m^{-2}

(or 5.67×10^{-8} J / m^2)

and T is the absolute temperature

(in Kelvin)

Energy = σT^4

Stefan-Boltzmann Law (easy way)

This law links:

(E) the total amount of energy flux that is emitted by a blackbody & (T) the body's temperature

(specifically, the 4th power of the body's absolute temperature)

"the hotter the body, the (much) greater the amount of energy flux or radiation"

p 33

The **total amount of energy** flux described by the Stefan-Boltzmann Law is proportional to the **area under the Planck function curve**

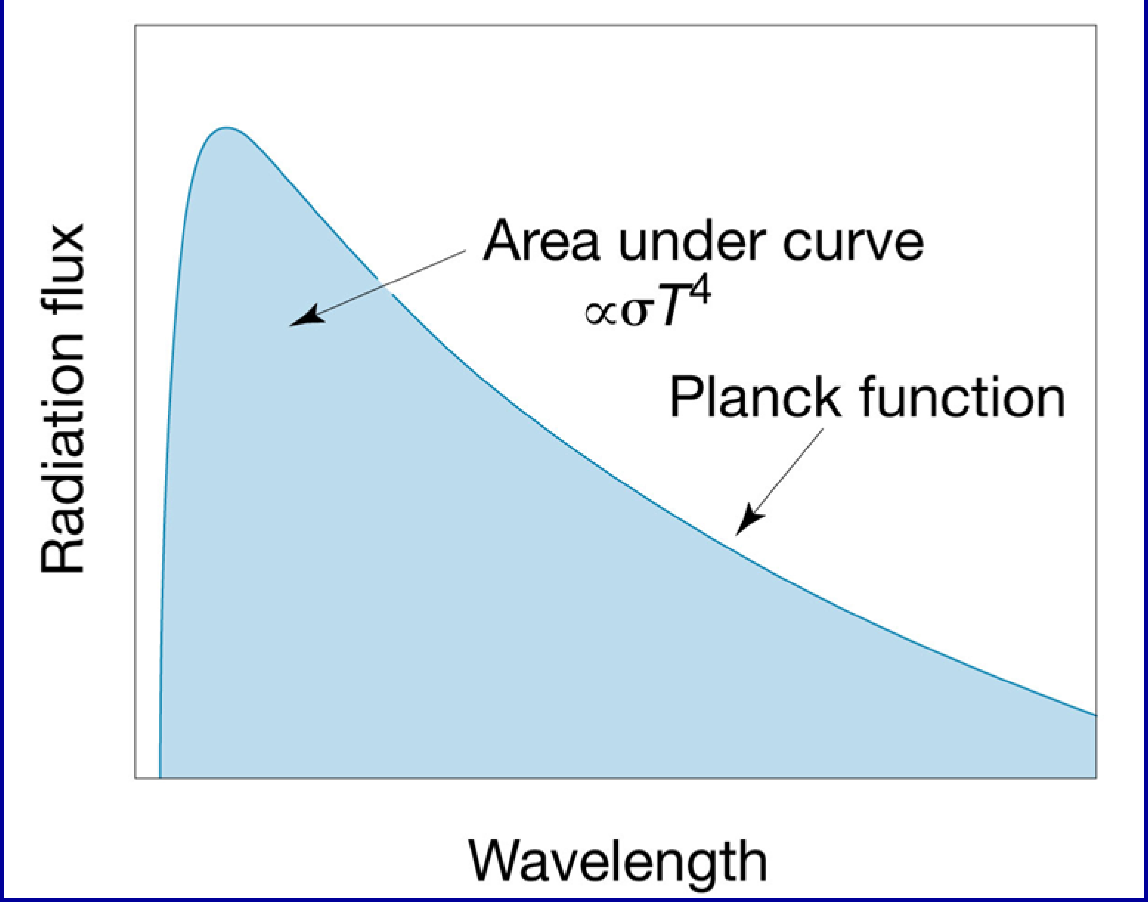


Figure on p 42 in SGC E-text

Stefan-Boltzmann Law:

"I'm HOT, so I emit LARGE amounts of high intensity energy"

"I'm COOL, so I emit LESSER amounts of energy. *plus* my ENERGY is at a lower intensity than Mr. Hotshot over there!"

Why is this concept important?

Because it means that:

the <u>amount</u> of radiation given off by a body is a <u>very *sensitive*</u> function of its <u>temperature</u>

Therefore . . . small changes in temperature can lead to BIG changes in the amount of radiation given off.

$$E = \sigma T^4$$

p 33

Q3 – Which would you use:
the Planck Function or the Stefan-Boltzmann Law
to accurately compute
the total amount of ENERGY
emitted to space by planet Earth?

1. The Planck Function

2. The Stefan Boltzmann Law

3. Both of them together

4. Neither one is appropriate because the Earth is NOT a blackbody

Q3 – Which would you use:
the Planck Function or the Stefan-Boltzmann Law
to accurately compute
the total amount of ENERGY
emitted to space by planet Earth?

1. The Planck Function
2. The Stefan Boltzmann Law
3. Both of them together
4. Neither one is appropriate because the Earth is NOT a blackbody

Q4 – Which would you use:
the Planck Function or the Stefan-Boltzmann Law
to compute the total amount of energy
emitted to space by planet Earth,
IF you assume the Earth emits like a blackbody
& you know the Earth's temperature?

1. The Planck Function

2. The Stefan Boltzmann Law

3. Neither one is appropriate because you would need to know the wavelengths of radiation the Earth emits

4. Don't know

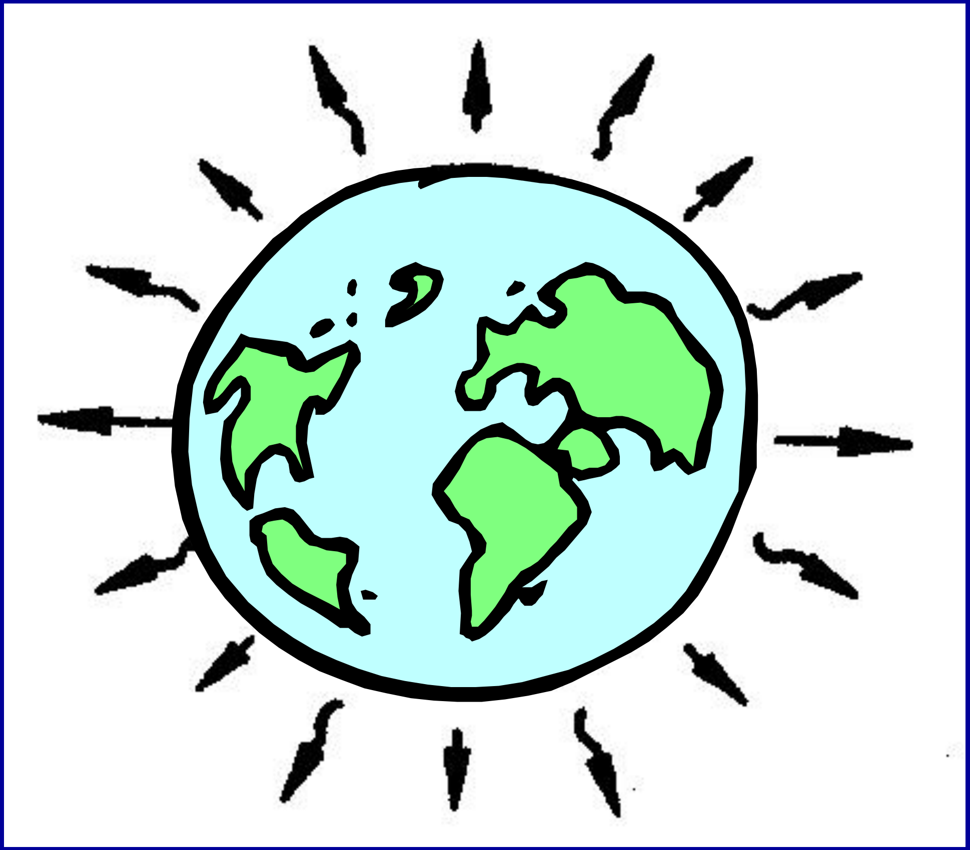
Q4 – Which would you use:
the Planck Function or the Stefan-Boltzmann Law
to compute the total amount of energy
emitted to space by planet Earth,
IF you assume the Earth emits like a blackbody
& you know the Earth's temperature?

1. The Planck Function
2. The Stefan Boltzmann Law
3. Neither one is appropriate because you would need to know the wavelengths of radiation the Earth emits
4. Don't know

How to do it:

$$E = \sigma T^4$$

E = **Energy per unit area**, so all we need to know is the **AREA** of the emitting Earth's surface + what T is.



From geometry: Do you remember the formula for computing the area of a sphere?

The area of a sphere of radius R is
$4 \pi R^2$

$$E = 4 \pi R^2 \times \sigma T^4$$

See box on p 44 in SGC E-text for more details

10 minute SUSTAINABILITY SEGMENT

more of:

Then . . . **MORE CLICKER PARTICIPATION POINTS**

LAW # 4: Temperature and wavelength

As substances get HOTTER, the wavelength at which radiation is emitted will become SHORTER.

This is called Wien's law.

Wien's Law can be represented as:

$$\lambda_m = a/T$$

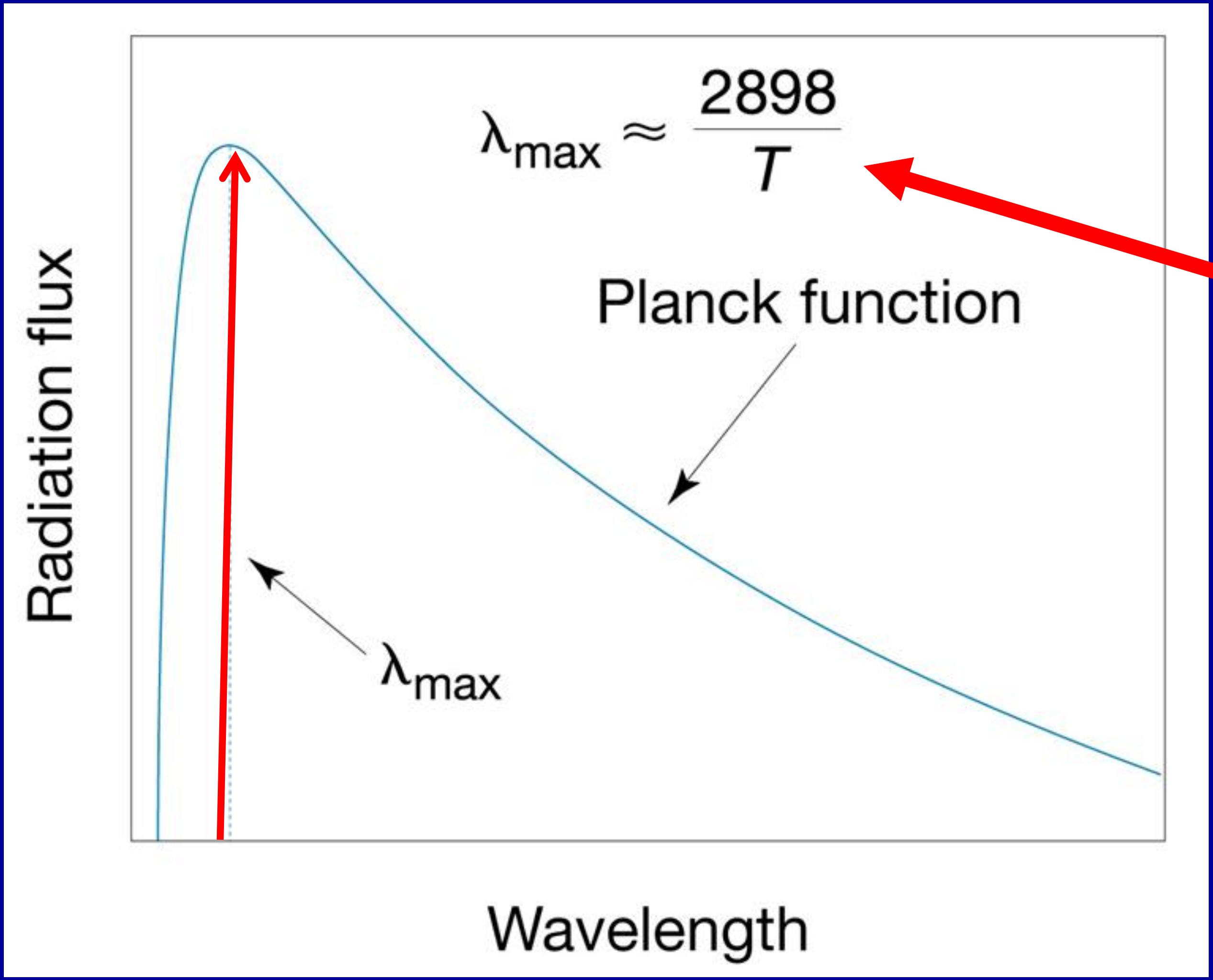
where λ_m is the **WAVELENGTH** in the spectrum at which the energy peak occurs,

(m indicates "max")

T is the absolute **TEMPERATURE** of the body, and

a is a **constant** (with a value of 2898)

(if λ_m is expressed in micrometers.)



Note the INVERSE relationship between wavelength and temperature

See figure on p 42 in SGC

Wien's Law (easy way)

$$\lambda_{max} = constant / T$$

(Inverse relationship between wavelength and temperature)

"The hotter the body, the shorter the wavelength"

"The cooler the body, the longer the wavelength"

Wien's Law -- Why is this concept important?

Because it means that very HOT objects (like the sun) that radiate like blackbodies will radiate the maximum amount of energy at SHORT wavelengths,

while COOLER bodies will radiate most of their energy at LONGER wavelengths.

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Wein's Law:

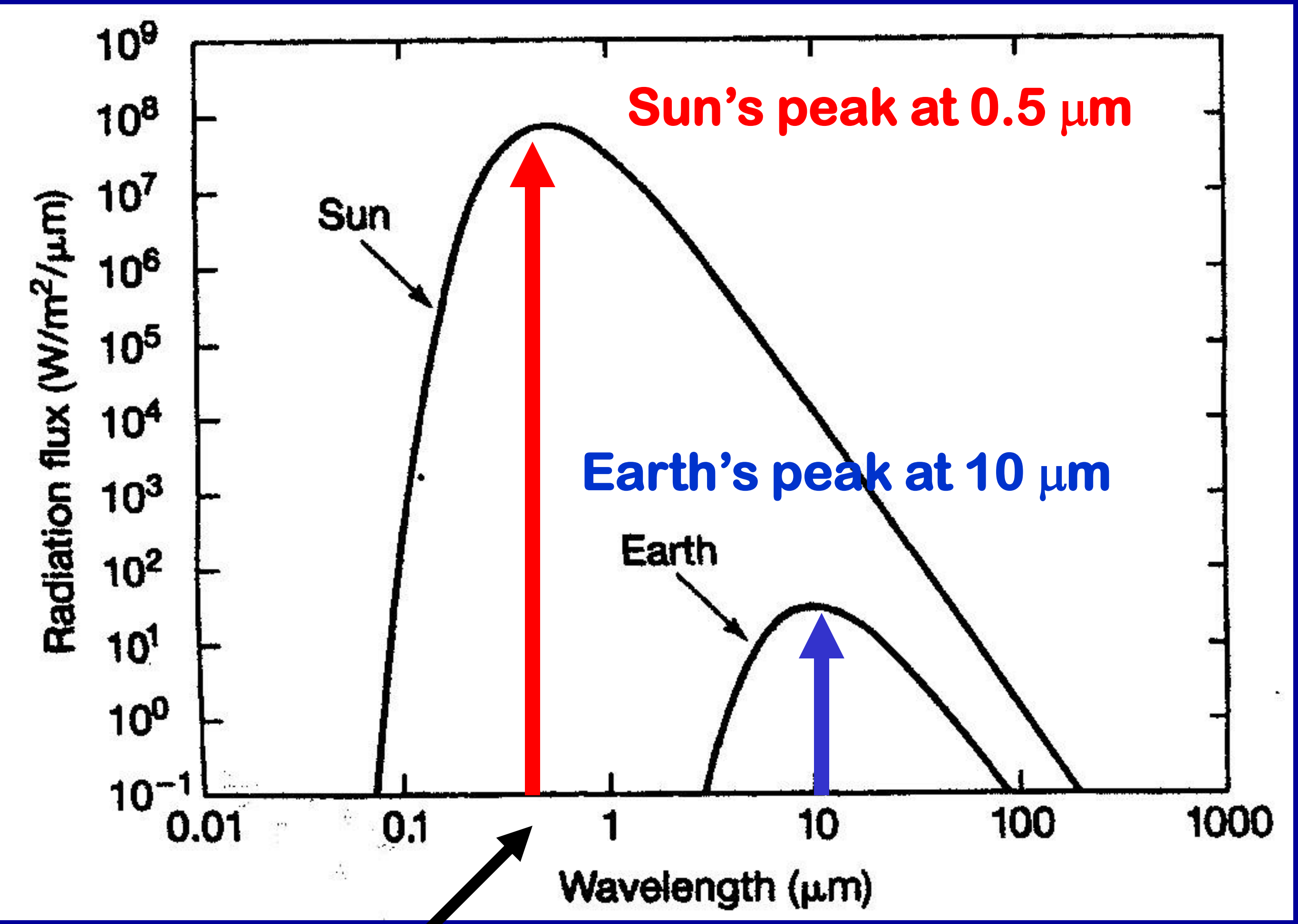
"I'm HOT, so I emit my maximum amount of radiation at SHORTER wavelengths"

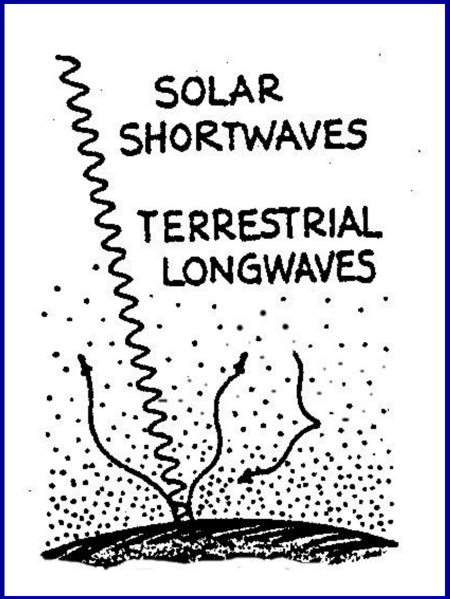
"I'm COOL, so I emit my maximum amount of radiation at LONGER wavelengths"

LW = infrared (IR)

SW = visible (VIS) & ultraviolet (UV)

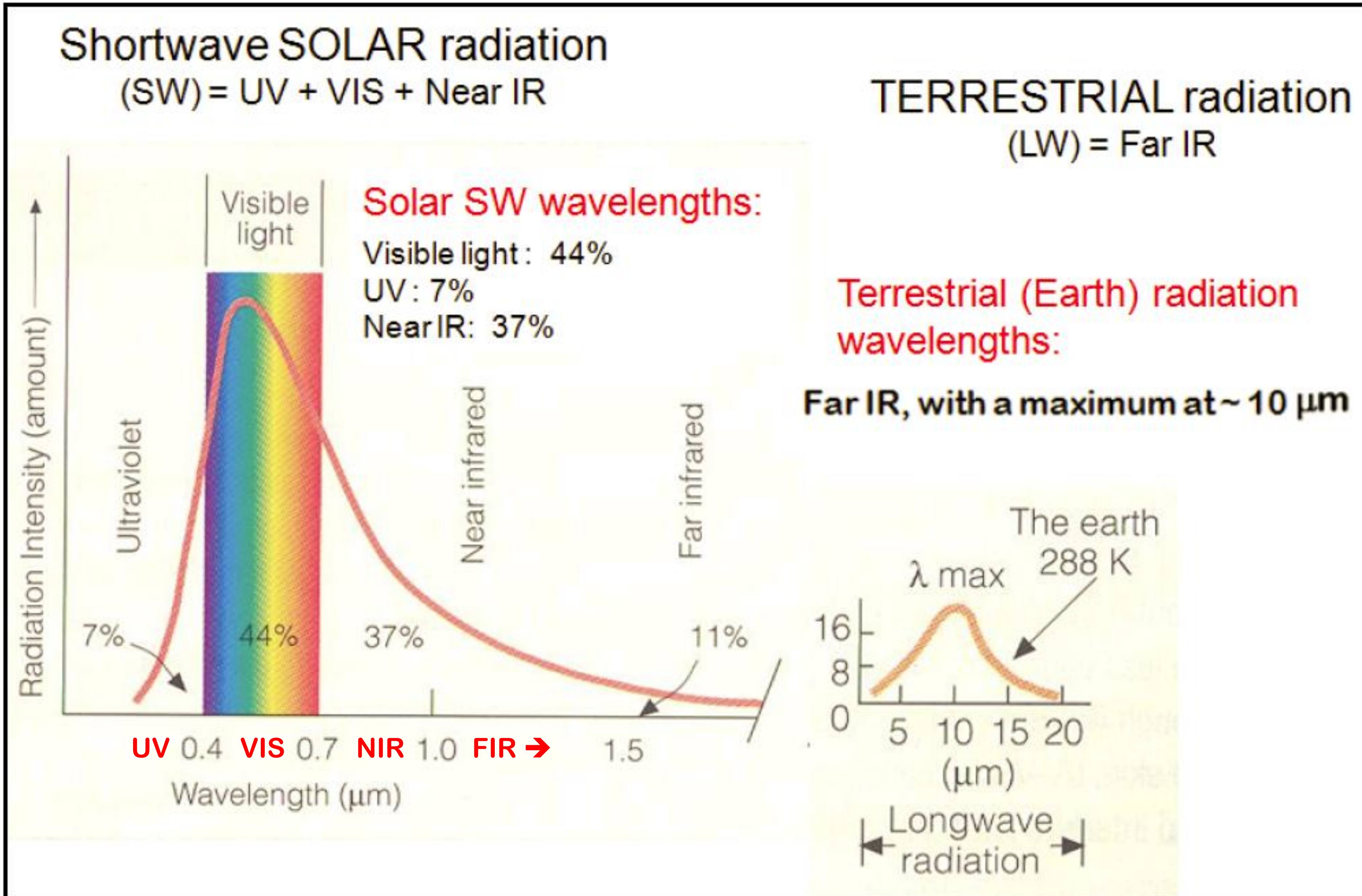
Wein's is the law behind this cartoon (on p 31)





NOTE: this is a logarithmic scale -- values increase exponentially to the right

Another view of the same concept:



Know & understand what this figure is illustrating!

REVIEW:

Q5 – Which choice correctly matches the Stefan-Boltzmann LAW with its "mantra" (A, B, C):

A "The hotter the body, the shorter the wavelength"
The cooler the body, the longer the wavelength"

B "SHORTER wavelengths have HIGHER intensity radiation than LONGER wavelengths "

C "The hotter the body, the (much) greater the amount of energy flux or radiation"

Top of p 35

(A) Wein's Law: $\lambda_m = a / T$

"The hotter the body, the shorter the wavelength"
The cooler the body, the longer the wavelength"

(B) Planck Function: $E = h c / \lambda$

"SHORTER wavelengths have HIGHER intensity radiation than LONGER wavelengths"

(C) Stefan-Boltzmann Law: $E = \sigma T^4$

"The hotter the body, the (much) greater the amount of energy flux or radiation"

☺